

The recent cosmological measurements vs $f(R)$ theory

Aleksander Łukasz Lenart

Astronomical Observatory of Jagiellonian University, Kraków

KMPS UJ

SeMPowisko 2024 19.04.2024

The simple description of Universe

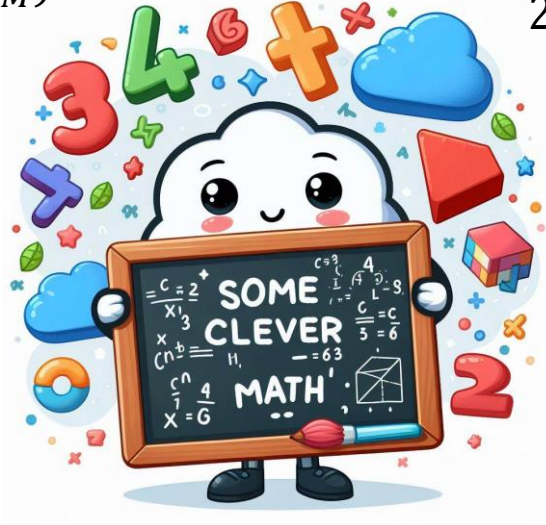
- Ω_M how much matter there is in the Universe
- Ω_Λ how much dark energy there is in the Universe
- $\Omega_k = 1 - \Omega_M - \Omega_\Lambda$ what is the curvature of the Universe
- H_0 what is the speed of expansion of the Universe
- z „redshift“, parameter characterising the observed electro-magnetic spectrum

$$d(z) = d(z, H_0, \Omega_M, \Omega_\Lambda, \Omega_k)$$

Was Einstein right?

The simplest formulation

$$S = \frac{1}{2\kappa} \int d^4x R \sqrt{-g} + \int d^4x L_M(g_{\mu\nu}, \psi_M)$$



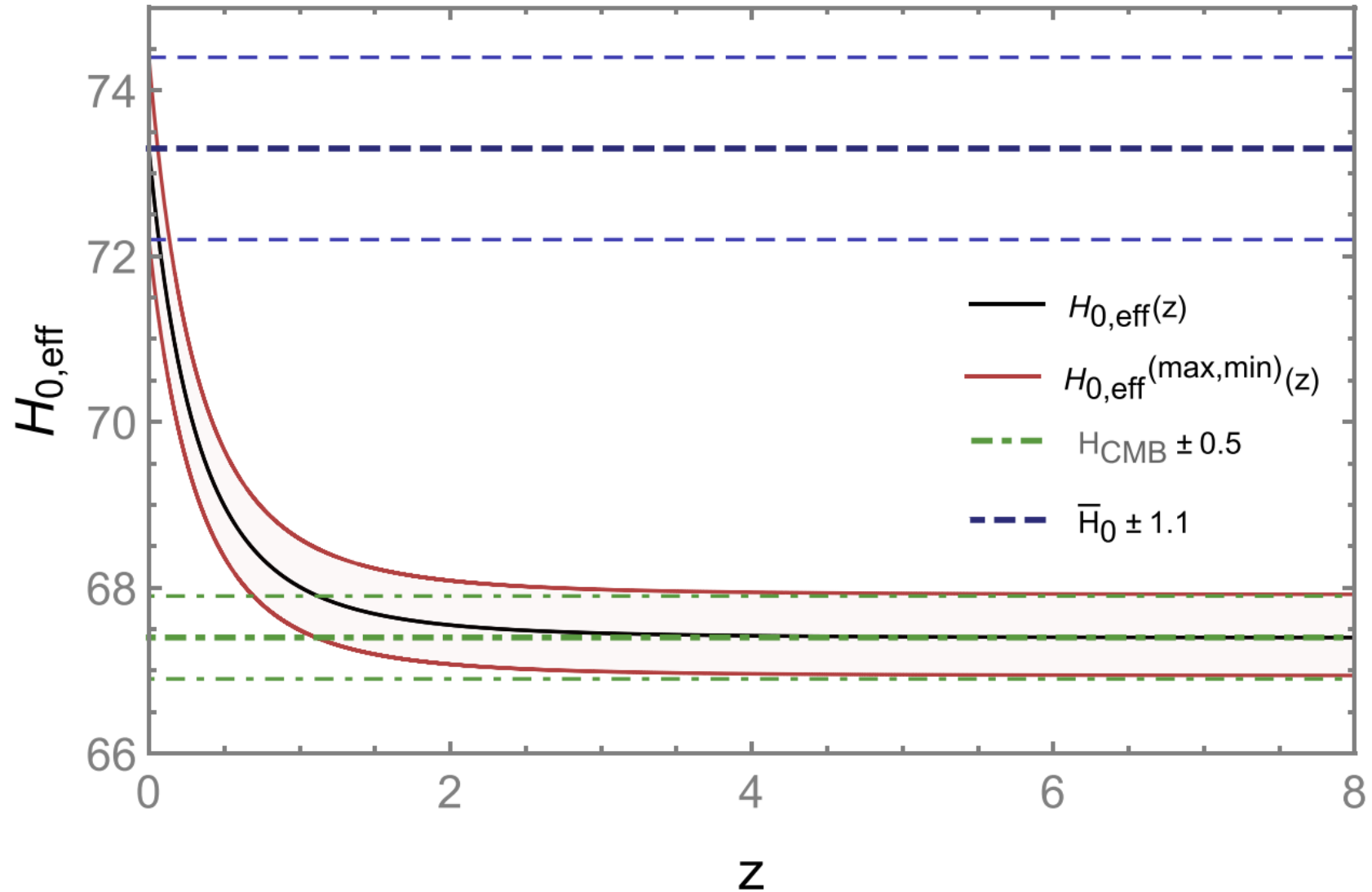
The $f(R)$ formulation

$$S = \frac{1}{2\kappa} \int d^4x f(R) \sqrt{-g} + \int d^4x L_M(g_{\mu\nu}, \psi_M)$$

$$d = (1+z) \frac{c}{H_0} \int_0^z \frac{d\zeta}{\sqrt{\Omega_M (1+\zeta)^3 + (1-\Omega_M)}}$$

$$d = (1+z) \frac{c}{H_0} \int_0^z \frac{d\zeta}{\sqrt{\frac{\Omega_M}{\varphi(\zeta)} (1+\zeta)^3 + (1-\Omega_M)}}$$

Impact on the measurements of a speed of expansion



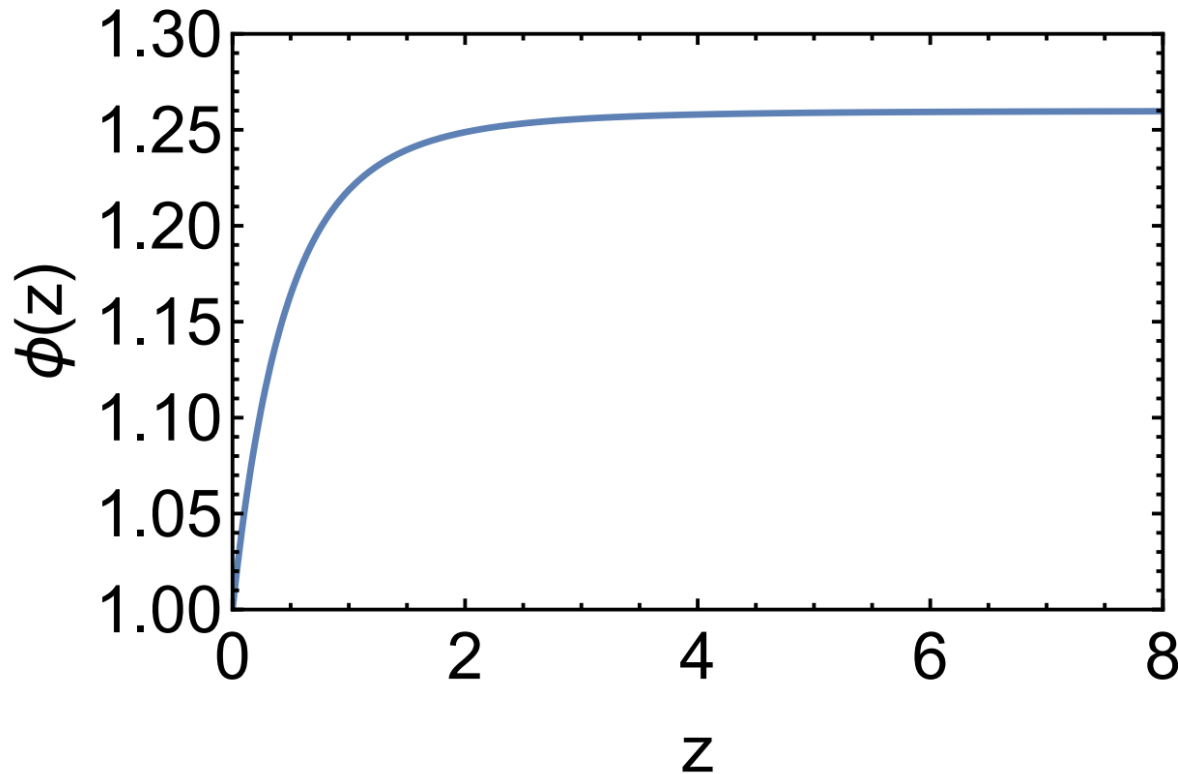
From equations of motion: $x \stackrel{\text{def}}{=} \ln(1+z)$

$$d = (1+z) \frac{c}{H_0} \int_0^z \frac{d\zeta}{\sqrt{\frac{\Omega_M}{\varphi(\zeta)} (1+\zeta)^3 + (1-\Omega_M)}}$$

$$2(\Omega_M - 1)\varphi(x)^2(\varphi'(x) + \varphi''(x)) + e^{3x}\Omega_M(\varphi'(x)^2 - \varphi(x)(7\varphi'(x) + 2\varphi''(x))) = 0$$

Initial values:

$$\varphi(0) = 1, \quad \varphi'(0) = w$$



$$\Omega_M = 0.30$$
$$\varphi'(x=0) = -0.56$$

But it should modify also Ω_M

$$d = (1+z) \frac{c}{H_0} \int_0^z \frac{d\zeta}{\sqrt{\Omega_M(1+\zeta)^3 + (1-\Omega_M)}}$$

$$d = (1+z) \frac{c}{H_0} \int_0^z \frac{d\zeta}{\sqrt{\frac{\Omega_M}{\varphi(\zeta)}(1+\zeta)^3 + (1-\Omega_M)}}$$

If we fit the simple model at high z , we expect smaller Ω_M , than at small z

Do we actually observe it?



Quasars ($z < 7.5$) best fit

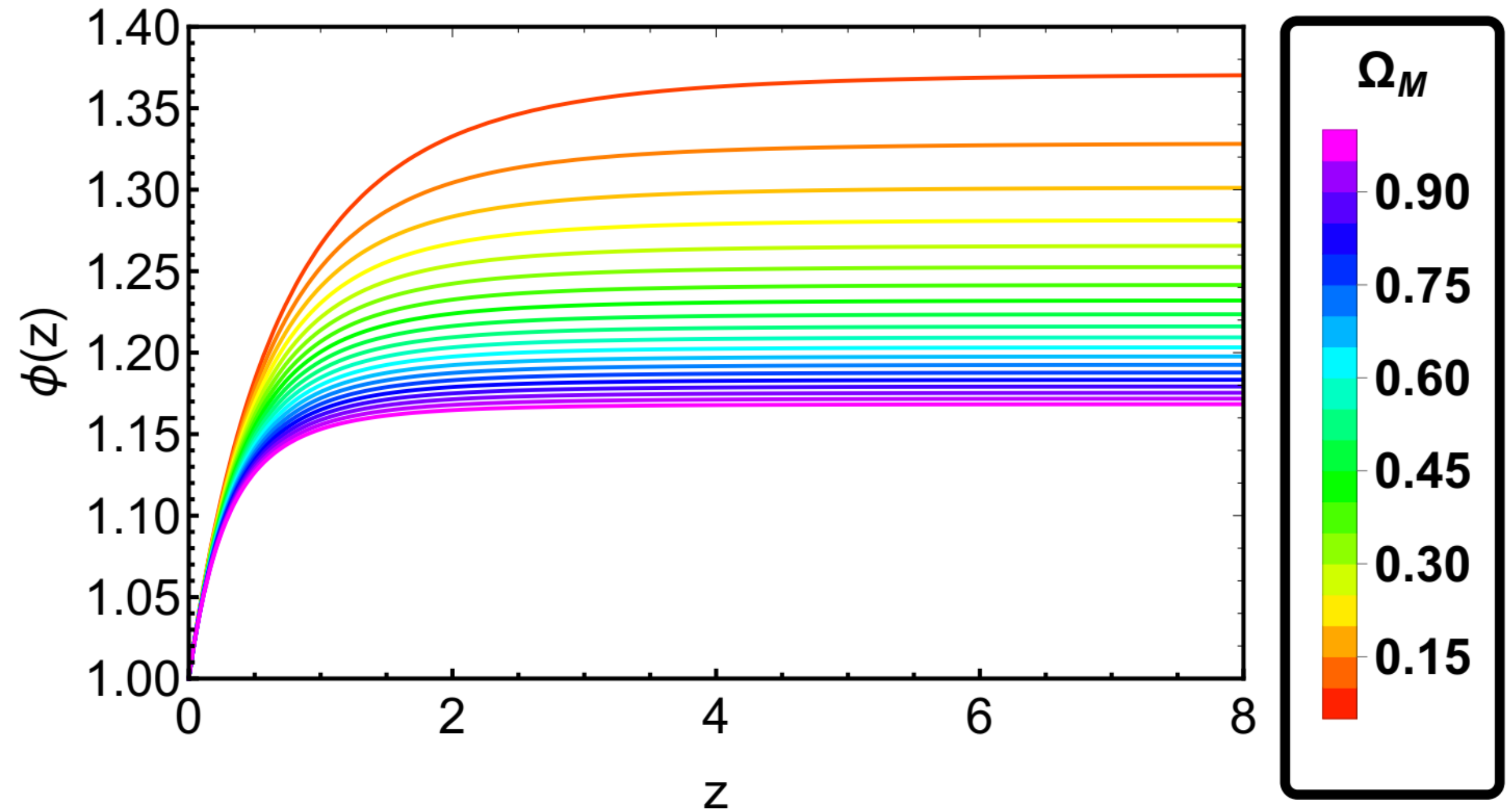
$$\Omega_M = 0.24 \pm 0.06$$

γ -ray attenuation ($z < 6$)
best fit

$$\Omega_M = 0.19 \pm 0.08$$

SNe Ia ($z < 2.1$) best fit

$$\Omega_M = 0.334 \pm 0.018$$



Thank You For Attention

Aleksander Łukasz Lenart

aleksander.lenart@student.uj.edu.pl

Astronomical Observatory of Jagiellonian University, Kraków

Sources

- De Felice, A., Tsujikawa, S. $f(R)$ Theories. Living Rev. Relativ. 13, 3 (2010). <https://doi.org/10.12942/lrr-2010-3>
- Dainotti M.G., Lenart A.ł., Yengejeh M.G., Chakraborty S., Fraija N., Di Valentino E., Montani G., 2024, PDU, 44, 101428. doi:10.1016/j.dark.2024.101428
- Montani G., De Angelis M., Bombacigno F., Carlevaro N., 2024, MNRAS, 527, L156. doi:10.1093/mnrasl/slad159
- Brout D., Scolnic D., Popovic B., Riess A.~G., Carr A., Zuntz J., Kessler R., et al., 2022, ApJ, 938, 110. doi:10.3847/1538-4357/ac8e04
- Domínguez A., Østergaard Kirkeberg P., Wojtak R., Saldana-Lopez A., Desai A., Primack J.R., Finke J., et al., 2024, MNRAS, 527, 4632. doi:10.1093/mnras/stad3425